

Two-level hierarchical fragmentation in the Orion Molecular Cloud 1 northern filament

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ABSTRACT

Context. The filamentary structure of molecular clouds may set important constraints on the mass distribution of stars forming within them. It is therefore important to understand what physical mechanism dominates filamentary cloud fragmentation and core formation. **Aims.** Orion A is the nearest giant molecular cloud, and its “ \int -shaped filament” is a very active star forming region that is a good target for such a study. We have recently reported on the collapse and fragmentation properties of the northernmost part of this structure, located ~ 2.4 pc north of Orion KL – the Orion Molecular Cloud 3 (OMC 3, Takahashi et al. 2013). As part of our project to study the \int -shaped filament, we analyze the fragmentation properties of the northern OMC 1 filament (located $\lesssim 0.3$ pc north of Orion KL). This filament is a dense structure previously identified by JCMT/SCUBA submillimeter continuum and VLA NH_3 observations and shown to have fragmented into clumps. Our aim is to search for cores and young protostars embedded within OMC 1n, and to study how the filament is fragmenting to form them.

Methods. We observed OMC 1North (hereafter OMC 1n) with the Submillimeter Array (SMA) at 1.3 mm and report on our analysis of the continuum data.

Results. We discovered 24 new compact sources, ranging in mass from 0.1 to 2.3, in size from 400 to 1300 au, and in density from 2.6×10^7 to $2.8 \times 10^6 \text{ cm}^{-3}$. The masses of these sources are similar to those of the SMA protostars in OMC 3 (Takahashi et al. 2013), but their typical sizes and densities are lower by a factor of ten. Only 8% of the new sources have infrared counterparts, yet there are five associated CO molecular outflows. These sources are thus likely in the Class 0 evolutionary phase yet it cannot be excluded that some of the sources might still be pre-stellar cores. The spatial analysis of the protostars shows that these are divided into small groups that coincide with previously identified JCMT/SCUBA 850 μm and VLA NH_3 clumps, and that these are separated by a quasi-equidistant length of $\approx 30'$ (0.06 pc). This separation is dominated by the Jeans length, and therefore indicates that the main physical process in the filament’s evolution was thermal fragmentation. Within the protostellar groups, the typical separation is $\approx 6''$ (~ 2500 au), which is a factor 2-3 smaller than the Jeans length of the parental clumps within which the protostars are embedded. These results point to a hierarchical (2-level) thermal fragmentation process of the OMC 1n filament.

Key words. Techniques: interferometric – Stars: formation – ISM: structure – ISM: clouds

1. Introduction

It is well established that filamentary molecular cloud structures are ubiquitous (Schneider & Elmegreen 1979). New observations have generated a revival in the star formation community’s interest in these structures, namely observations from the Herschel Space Observatory (e.g. André et al. 2010) that revealed the filamentary nature of clouds in great detail. Combined with a better characterization of the spatial structure of clouds, there has also been advancement in the understanding of the kinematics of filamentary structures. Of particular note is the work carried out by Hacar et al. (2013) where the Taurus filament L1495/B213 was found to have undergone hierarchical fragmentation and is composed of several velocity coherent sub-filaments, each of which has sonic internal velocity dispersion and with a mass-per-unit length consistent with an isothermal cylinder at 10 K. Hierarchical thermal fragmentation has also been inferred to have occurred in other star forming regions, e.g., in the Orion complex (Takahashi et al. 2013) and in the Spokes cluster in NGC 2264 (Teixeira et al. 2006, 2007;

Pineda & Teixeira 2013). These are much denser and richer star forming environments than Taurus, yet the spatial distribution of protostars within filaments are consistent with the same fragmentation process.

The Orion Molecular Cloud (OMC) A is the nearest (414 ± 7 pc, Menten et al. 2007) and one of the richest star-forming giant molecular clouds. It is an elongated structure (“ \int -shaped filament”, Bally et al. 1987) spanning ~ 10 pc along the north-south direction (e.g. Johnstone & Bally 1999), comprising of several large clouds to the north (OMC 1, OMC 2, and OMC 3) and south (OMC 4 and OMC 5). These clouds retain the overall filamentary nature of the region and have been the subject of numerous observational studies (see the reviews Muench et al. 2008; O’Dell et al. 2008; Peterson & Megeath 2008, and references within).

Of particular interest is the OMC 1 cloud for it is the most massive ($> 2200 M_\odot$) component of the OMC (Bally et al. 1987) and has spawned the Orion Nebula Cluster (ONC; O’Dell 2001), which is presently located in front of the cloud. As shown in Figure 1, the OMC 1 cloud is comprised of dense filaments in

the north (OMC 1n) that connect to Orion KL, and of the very active star forming region Orion South (e.g. Zapata et al. 2004). Wiseman & Ho (1998) carried out interferometric NH_3 observations of the OMC 1n filaments with the VLA using high spectral (0.3 km s^{-1}) and moderate angular ($8''$) resolution; their observations identified filamentary structures throughout a 0.5 pc region extending to the north from Orion-KL. These northern filaments, henceforth referred to as OMC 1n, were found to be fragmented into bead-like chains of dense clumps. We re-observed these filaments with the Submillimeter Array (SMA) in order to further characterize their fragmentation and the dense NH_3 clumps previously identified. We found 24 previously unknown compact millimeter sources, and their spatially distribution indicates fragmentation of the filament at two length scales. This paper is part of a series of high-angular resolution OMC studies focused on the filamentary structures and their embedded sources, the first of which is Takahashi et al. (2013) where we discuss the global fragmentation properties of the OMC and of the OMC 3 cloud in more detail.

We describe our SMA observations of OMC 1n in § 2 and present our results in § 3. Finally, we discuss our findings on the characterization of these new millimeter sources in § 4 and summarize our conclusions in § 5

2. Observations and Data Reduction

We observed OMC 1n using the eight 6 m antennas of the SMA¹ (Ho et al. 2004) in the compact configuration on December 4, 20, and 24, 2008 (2007B-S028, P.I.: L. Zapata) and in the subcompact configuration on February 1, 2009 (2008B-S072, P.I.: L. Zapata) at 230 GHz. The primary beam at this frequency is $55''$, and the nine pointing centers were distributed in a Nyquist sampled grid. Figure 1 shows the area surveyed by the SMA where the white circles represent the primary beam size pointings or fields. We observed each field for 1 minute during each of 30 cycles through all the fields, giving a total on-source integration time of 30 minutes. Both the LSB ($\nu_c=230 \text{ GHz}$) and USB ($\nu_c=220 \text{ GHz}$) data were obtained simultaneously with the 90° phase switching technique by the digital spectral correlator, which had a bandwidth of 2 GHz for each sideband.

The SMA correlator covers 2 GHz bandwidth in each of the two side-bands separated by 10 GHz. Each band is divided into 24 chunks of 104 MHz width, which were covered by fine spectral resolution (256 channels correspond to a velocity resolution of 0.53 km s^{-1} in our observing setting). In addition to continuum observations, molecular lines such as CO (2–1), ^{13}CO (2–1), C^{18}O (2–1), and ^{13}CS (5–4), were simultaneously obtained with the 2 GHz bandwidth in each sideband. These molecular line data will be presented in a future article.

The continuum data from both sidebands (USB + LSB) were combined to obtain a higher signal-to-noise ratio. The combined configurations of the arrays provided projected baselines ranging from 5.1 to $56.0 \text{ k}\lambda$. Our observations were insensitive to structures larger than $32''$ ($13000 \text{ AU}/0.07 \text{ pc}$) at the 10% level (Wilner & Welch 1994). Passband calibration was achieved by observing the quasars, 3C 454.3 and 3C 273. Amplitudes and phases were calibrated by observations of 0530+135 and 0541-056 for the compact configurations and 0607-085 and 0541-056 for the subcompact configurations. Those flux densities were determined relative to Uranus. The overall flux un-

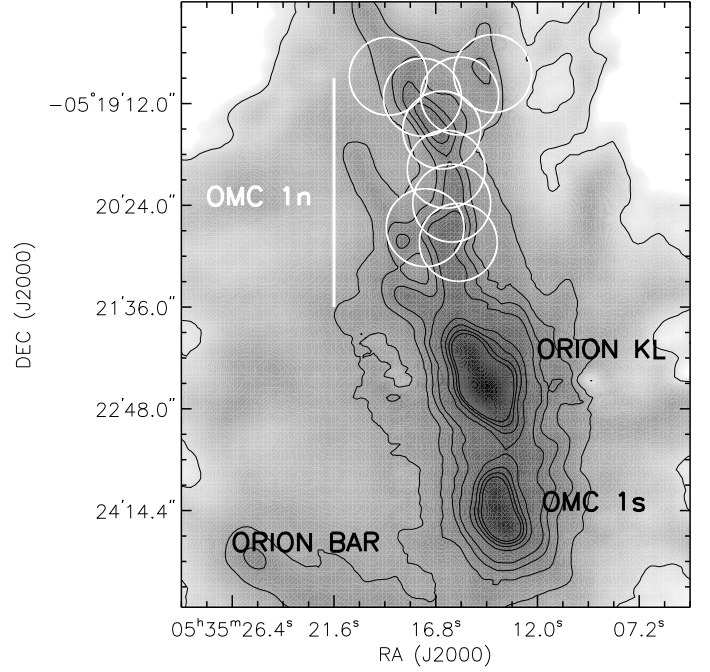


Fig. 1. JCMT/SCUBA $850 \mu\text{m}$ continuum image of the OMC-1 cloud (Johnstone & Bally 1999), with various features indicated as references, such as the Orion Bar, Orion South and Orion KL. The black contours range from 0.5 to 8.5 Jy beam^{-1} in steps of 2 Jy beam^{-1} , and from 15 to 30 Jy beam^{-1} in steps of 5 Jy beam^{-1} . The white circles represent the pointings made to cover the SMA area surveyed and their size correspond to the SMA primary beam, $55''$.

certainty was estimated to be $\sim 20\%$. The pointing accuracy of the SMA observations was $\sim 3''$. The raw data were calibrated using MIR, originally developed for the Owens Valley Radio Observatory (Scoville et al. 1993) and adopted for the SMA.

In order to produce the mosaic continuum image, certain frequencies (which include aforementioned molecular lines) are subtracted from the visibility data using a miriad task “uvlin”. The effective bandwidth for the continuum emission is approximately 3.6 GHz . We used the task CLEAN in miriad with the ROBUST parameter equal to 2, and then we combined the resulting images linearly with the task LINMOS. The final synthesized beam size results directly of combining all observations, and is $\approx 3''$ FWHM (corresponding to 1200 AU). The achieved mean rms noise level for the entire mosaic image is estimated to be $\approx 11 \text{ mJy}$. The SMA observational log for each date and the observational parameters are summarized in Table 1 and Table 2.

Table 1. SMA Observational Logs

Date	Configuration	Number of Antennae	τ_{225}
2007 Dec. 4	compact	7	0.05–0.08
2007 Dec. 20	compact	5	0.22–0.33
2007 Dec. 24	compact	7	0.12–0.24
2009 Feb. 1	subcompact	8	0.25–0.32

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Table 2. SMA Observational Parameters

Parameter	Value
Configurations	$3 \times$ compact and $1 \times$ subcompact
Primary beam HPBW (arcsec)	55
Grid spacing (arcsec)	25
Synthesized Beam HPBW (arcsec)	3.0
Equivalent frequency (GHz)	225
Total continuum bandwidth (GHz)	3.6
Projected base line range (k λ)	5.1 – 56.0
Maximum detectable structure (arcsec) ^a	32
Gain calibrators	0530+135, 0541-056 and 0607-085
Bandpass calibrator	3C 454.3 and 3C 273
Flux calibrators	Uranus
RMS noise level (mJy beam ⁻¹)	~11

Notes. ^(a) Our observations were insensitive to more extended emission than this size-scale structure at the 10% level (Wilner & Welch 1994).

3. Results

3.1. New compact continuum sources

We detected 24 previously unknown compact millimeter sources within the OMC-1N region; these new sources were identified as having peak fluxes $\geq 5\sigma$ (i.e. $\geq 55 \text{ mJy beam}^{-1}$) and are shown in Figure 2. There is fainter extended emission at the 3σ level that may correspond to smaller cores. Only two of these sources were found to have (nir- and mid) infrared counterparts; one of them coincides with SMA-10, and the other is located very near SMA-15, more precisely, with the small protrusion extending from the peak of the emission. We also detected red- and/or blue-shifted CO emission arising from bipolar outflows.

We used MIRIAD (Sault et al. 1995) to measure the fluxes, sizes and positions of each source (IMFIT procedure), where each source was modeled by an elliptical 2D Gaussian. As can be seen in Figure 2, the sources appear to be distributed in groups, i.e. they are connected as a single structure when using a 3σ contour. We therefore carried out multi-gaussian fits so that all sources in a particular group were fit together. There are six groups of sources that were fit simultaneously: group 1 consists of sources SMA-1 and SMA-2, group 2 consists of sources SMA-4 and SMA-5, group 3 consists of sources SMA-6, SMA-7, SMA-8, SMA-9 and SMA-10, group 4 consists of sources SMA-11, SMA-12, SMA-13 and SMA-14, group 5 consists of sources SMA-17, SMA-18 and SMA-19, and group 6 consists of sources SMA-21 and SMA-22. Figure 2 shows that SMA-15 may in fact consist of two sources (the presence of an infrared counterpart offset from the position of the main peak and coincident with the small south-west extended emission supports this), however we were unable to fit SMA-15 with two simultaneous 2D gaussians. Likewise for source SMA-20. The measured properties for each source are listed in Table 3. The total integrated flux listed was measured from the multi-2D-gaussian fitting. Sources SMA-9, SMA-11, and SMA-18 are not well fit by a 2D gaussian model, however it was necessary to include them in order to fit well the nearby brighter compact source(s) within the same group. The errors cited for the total flux correspond to the statistical uncertainties, however they are in fact dominated by the calibration uncertainty that is $\sim 20\%$, as previously mentioned in § 2. The total flux listed in Table 3 corresponds to thermal dust emission but may also have some contribution from thermal free-free emission from protostellar jets.

In order to calculate this latter contribution, we use a spectral index of 0.6 (Panagia & Felli 1975; Anglada et al. 1998) and estimate the flux from thermal free-free emission, S_ν , for these sources from $S_\nu \propto \nu^{0.6}$, where ν is the frequency of the emission. The SMA sources detected in OMC 1n share very similar properties with those in OMC 3 (Takahashi et al. 2013) and OMC 2 (Takahashi et al. 2015). We can therefore make use of VLA measurements at 3.6 cm of OMC 2 protostars (Reipurth et al. 1999, Table 1) as a benchmark, since most of the flux at this wavelength is from free-free emission. Using their minimum and maximum flux densities, we find that the contribution of thermal free-free emission from jets would range between 1.1 mJy and 20.9 mJy at 1.3 mm for the OMC 2 sources. We assume that the sources in OMC 1n would have a similar contribution, however, future observations are necessary to measure the exact amount of free-free emission for each of the OMC 1n sources. We note that the estimated flux from thermal free-free emission is typically within the error of the total integrated flux for each source.

Before proceeding, we would like to make a brief remark on the nomenclature used in this paper. The terms "core" and "clump" do not have a strict definition in the field of star formation and are sometimes used interchangeably in the literature, which may lead to erroneous comparisons of results. We use these two terms in the following manner: a core (or several cores) is formed within a clump and is therefore relatively denser than the clump, and we follow the nomenclature used in Takahashi et al. (2013), where "small-scale clumps" in the \int -shaped-filament are structures with sizes on the order of 0.3-0.1 pc, and "cores" are structures with sizes typically 10 times smaller. The aforementioned groups of sources would therefore correspond to small-scale clumps that fragmented into cores, e.g. group 1 corresponds to a single small-scale clump that fragmented into two cores, group 3 corresponds to a single small-scale clump that fragmented into five cores, etc. Each core corresponds to an SMA source.

The left panel of Figure 3 compares the SMA 1.3 mm sources (grayscale) with previously observed JCMT/SCUBA 850 μm emission (Johnstone & Bally 1999) (contours). Our SMA observations spatially resolved all the JCMT/SCUBA clumps, i.e. all of the single-dish identified clumps are associated with at least one SMA source. The right panel of Figure 3 compares the SMA 1.3 mm sources (grayscale) with previously observed VLA NH₃ (1,1) emission (Wiseman & Ho 1998) (contours). The SMA sources clearly lie within dense filamentary regions, although they are not always located at the center of an ammonia clump. For example, sources SMA-1 and SMA-2 are located at the center of the same 850 μm clump, yet are spatially offset by $\sim 10''$ from the center of an ammonia core. Other sources that are not located near the center of ammonia small-scale clumps are SMA-3, SMA-4, SMA-5, SMA-13, SMA-16, SMA-21, SMA-23, and SMA-24.

3.2. Masses and sizes

We can determine the masses of these sources if we assume the 1.3 mm emission arises from optically thin dust that can be characterized by a single temperature (e.g. Mundy et al. 1995; Bally et al. 1998). The total gas+dust mass is calculated from the following equation:

$$M = \frac{F_\nu d^2}{B_\nu(T_d) \kappa_\nu}, \quad (1)$$

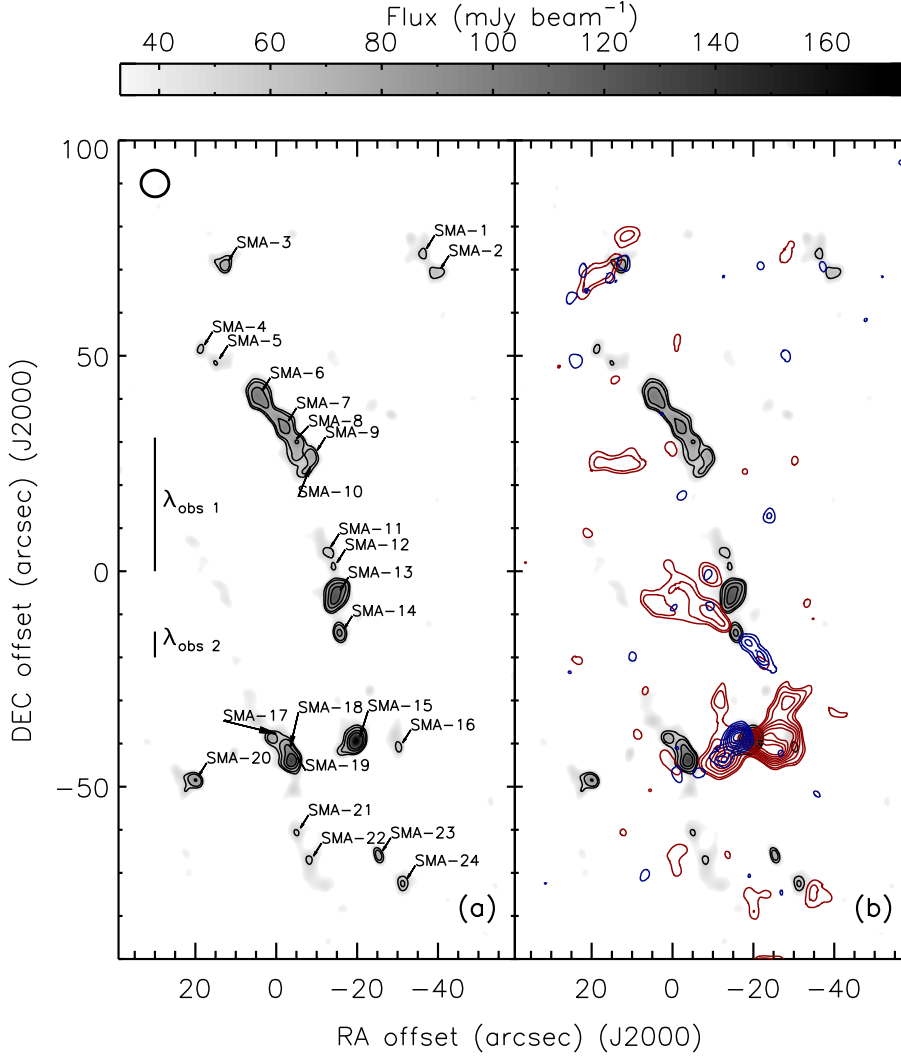


Fig. 2. SMA 1.3 mm continuum emission (grayscale, lowest level corresponds to $3\text{-}\sigma$ emission), where the black contours range from 5σ to 10σ in steps of 1σ ($\sigma \sim 11 \text{ mJy beam}^{-1}$). Coordinate offsets are measured with respect to $(\alpha, \delta) = (05^{\text{h}}35^{\text{m}}17.0^{\text{s}}, -05^{\circ}22'00.0'')$ (J2000.0). In panel (a), the new sources are indicated by arrows with their corresponding IDs (compare with Tables 3 and 4). The vertical scale bars correspond to the median 1st and 7th nearest neighbor separations of the SMA sources, $\lambda_{\text{obs } 1}$ and $\lambda_{\text{obs } 2}$, respectively (see Section 3.3 and Table 5). Panel (b) shows CO (2-1) red- and blue-shifted emission from molecular outflows, in red and blue contours, respectively. The red contour levels range from 14 to $136 \text{ Jy beam}^{-1} \text{ km s}^{-1}$, in steps of $14 \text{ Jy beam}^{-1} \text{ km s}^{-1}$. The blue contour levels range from 9 to $31 \text{ Jy beam}^{-1} \text{ km s}^{-1}$, in steps of $3 \text{ Jy beam}^{-1} \text{ km s}^{-1}$. The molecular outflow velocities range from 0 to 40 (red-shifted emission), and from -20 to 0 (blue-shifted emission).

where F_{ν} is the measured total flux of the source, d is the distance to the source, $B_{\nu}(T_d)$ is the Planck function for a dust temperature T_d , and κ_{ν} is the dust mass opacity. We follow Menten et al. (2007) and adopt a distance of $414 \pm 7 \text{ pc}$ to the SMA sources. The dust mass opacity is calculated from $\kappa_{\nu} = 0.1(\nu/1000 \text{ GHz})^{\beta} \text{ cm}^2 \text{ g}^{-1}$ (Beckwith et al. 1990), assuming a gas-to-dust ratio of 100, and where we take the emissivity index, β , to be 1.5, following Johnstone & Bally (1999) and the latest results from Planck for this region (Lombardi et al. 2014); κ_{ν} is therefore $\kappa_{230 \text{ GHz}} = 11 \times 10^{-3} \text{ cm}^2 \text{ g}^{-1}$. According to Wiseman & Ho (1998), the filament within which the SMA sources are embedded in has a kinetic gas temperature of 15 K, and this is the temperature we assume for T_d in our mass calculation (assuming that the gas and dust temperatures are coupled). Using the integrated flux densities given in Table 3, we calculate the masses for each source and their values are shown in Table 4. We calculated the radius of each source by taking the geometric mean of the deconvolved major and minor axes (given in Table 3), $R_{\text{geom}} = \sqrt{\text{size}_{\text{maj. axis}} \cdot \text{size}_{\text{min. axis}}}$, which was in turn used to determine the number densities, \bar{n}_{H_2} , for each source assuming spherical symmetry. Both R_{geom} and \bar{n}_{H_2} are shown in Table 4.

Tatematsu et al. (2008) find a constant temperature of $\sim 20 \text{ K}$ over the entire \int -shaped filament. Here we opted for using the temperature measured by Wiseman & Ho (1998) because this value was obtained from higher angular resolution in-

terferometric observations. The single-dish observations from Tatematsu et al. (2008) include extended emission of the cloud that envelope the small-scale clumps and that is presumably warmer, thus providing a slightly higher temperature. This effect is also suggested by Herschel data (e.g. see Fig. 2 of Lombardi et al. 2014). On the other hand, the interferometric observations filtered out the extended emission and should give a more reliable estimate of the temperature of the dense regions within the filament. For comparison purposes we added a note to Table 4 where we show that changing the temperature assumption to 20 K would lower the masses by factor of ~ 1.5 .

3.3. Spatial distribution

Figure 3 shows that the SMA sources are distributed along dense filamentary structures, and part of groups that show a quasi-equal spacing between (and within) them. To better describe quantitatively these spatial scales associated with the SMA sources, we present in Figure 4 an analysis of the 1st to the 10th (projected) nearest neighbor separations (NNS). Panel (a) shows the measured median value of the Nth ($N = 1, \dots, 10$) NNS in black vs. their σ -value. The plot shows that the empirical function has two minima, namely, for the 1st and 7th nearest neighbors. A minima in the $\sigma_{\text{Nth NNS}}$ function indicates a characteristic spatial scale of the distribution of sources (i.e., the width of the NNS distribution narrows as sources have similar sepa-

Table 3. Measured observables of the SMA sources from 2D elliptical Gaussian fitting.

ID	R.A. (J2000.0)	Dec. (J2000.0)	Pos. unc. (arcsec)	Size (arcsec)	P.A. (deg.)	$F_{230\text{ GHz}}^{\text{peak}}$ (mJy beam ⁻¹)	$F_{230\text{ GHz}}^{\text{int.}}$ ^a (mJy)
SMA-1 ^c	05:35:14.59	-05:18:46.1	0.3	$(3.5 \pm 1.0) \times (2.9 \pm 0.9)$	-21 ± 68	60 ± 3	118 ± 50
SMA-2 ^c	05:35:14.35	-05:18:50.8	0.3	$(6.9 \pm 1.7) \times (3.5 \pm 0.8)$	-79 ± 17	64 ± 2	224 ± 74
SMA-3	05:35:17.85	-05:18:48.9	0.2	$(4.9 \pm 1.3) \times (4.3 \pm 1.2)$	-33 ± 102	76 ± 5	227 ± 88
SMA-4 ^d	05:35:18.25	-05:19:08.4	0.3	$(4.5 \pm 1.7) \times (1.6 \pm 0.4)$	-23 ± 23	60 ± 4	112 ± 52
SMA-5 ^d	05:35:18.00	-05:19:11.7	0.3	$(2.7 \pm 1.5) \times (1.6 \pm 0.9)$	22 ± 100	58 ± 4	83 ± 66
SMA-6 ^e	05:35:17.29	-05:19:19.1	0.2	$(7.5 \pm 1.0) \times (5.2 \pm 0.4)$	23 ± 12	104 ± 6	485 ± 77
SMA-7 ^e	05:35:16.87	-05:19:26.1	0.2	$(7.1 \pm 2.3) \times (4.6 \pm 0.4)$	36 ± 13	95 ± 7	389 ± 137
SMA-8 ^e	05:35:16.67	-05:19:32.3	0.3	$(7.8 \pm 3.4) \times (4.4 \pm 1.1)$	-51 ± 28	66 ± 22	290 ± 174
SMA-9 ^{b,e}	05:35:16.39	-05:19:32.9	0.4	$(3.2 \pm 1.8) \times (2.6 \pm 1.6)$	4 ± 71	39 ± 25	30 ± 31
SMA-10 ^e	05:35:16.50	-05:19:36.7	0.3	$(8.0 \pm 3.4) \times (2.4 \pm 0.9)$	-53 ± 16	59 ± 19	197 ± 129
SMA-11 ^f	05:35:16.17	-05:19:55.3	0.3	$(5.6 \pm 1.5) \times (4.4 \pm 1.4)$	74 ± 133	61 ± 7	204 ± 86
SMA-12 ^{b,f}	05:35:16.06	-05:19:59.0	0.5	$(3.1 \pm 2.1) \times (2.4 \pm 1.8)$	-29 ± 70	32 ± 27	22 ± 30
SMA-13 ^f	05:35:16.01	-05:20:05.4	0.1	$(7.0 \pm 0.4) \times (4.4 \pm 0.2)$	-30 ± 6	128 ± 5	509 ± 44
SMA-14 ^f	05:35:15.95	-05:20:14.4	0.2	$(3.6 \pm 0.4) \times (2.6 \pm 0.3)$	5 ± 20	91 ± 7	171 ± 33
SMA-15	05:35:15.69	-05:20:39.4	0.1	$(3.3 \pm 0.1) \times (2.3 \pm 0.1)$	-32. ± 6	173 ± 2	301 ± 14
SMA-16	05:35:14.98	-05:20:40.8	0.3	$(5.5 \pm 1.6) \times (2.4 \pm 0.6)$	14 ± 17	61 ± 1	146 ± 55
SMA-17 ^g	05:35:17.07	-05:20:38.7	0.2	$(3.6 \pm 0.9) \times (1.9 \pm 0.3)$	-13 ± 16	74 ± 6	213 ± 66
SMA-18 ^g	05:35:16.84	-05:20:41.8	0.3	$(10.3 \pm 13.7) \times (1.1 \pm 0.6)$	-20 ± 18	55 ± 70	186 ± 358
SMA-19 ^g	05:35:16.72	-05:20:44.2	0.1	$(5.2 \pm 1.9) \times (3.8 \pm 2.7)$	63 ± 64	118 ± 37	199 ± 171
SMA-20	05:35:18.35	-05:20:48.5	0.2	$(4.7 \pm 0.8) \times (2.9 \pm 0.5)$	75 ± 25	86 ± 5	207 ± 49
SMA-21 ^h	05:35:16.67	-05:21:00.6	0.3	$(5.8 \pm 2.6) \times (3.5 \pm 1.6)$	-2 ± 56	57 ± 1	170 ± 107
SMA-22 ^h	05:35:16.45	-05:21:07.1	0.3	$(4.1 \pm 1.1) \times (3.7 \pm 1.3)$	-13 ± 99	58 ± 2	141 ± 63
SMA-23 ^b	05:35:15.31	-05:21:05.9	0.2	$(5.3 \pm 0.6) \times (3.2 \pm 0.3)$	24 ± 7	76 ± 4	120 ± 19
SMA-24	05:35:14.91	-05:21:12.3	0.3	$(5.3 \pm 0.7) \times (4.2 \pm 0.5)$	-4 ± 21	66 ± 3	203 ± 36

Notes. Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds.

(^a) Total integrated flux. We estimate the thermal free-free contribution to range between 1.1 mJy and 20.9 mJy (see text for details). This value is typically smaller than the error of $F_{230\text{ GHz}}^{\text{int.}}$.

(^b) un-deconvolved.

(^c) simultaneous fit of the group of sources SMA-1 and SMA-2.

(^d) simultaneous fit of the group of sources SMA-4 and SMA-5.

(^e) simultaneous fit of the group of sources SMA-6, SMA-7, SMA-8, SMA-9, SMA-10.

(^f) simultaneous fit of the group of sources SMA-11, SMA-12, SMA-13, and SMA-14.

(^g) simultaneous fit of the group of sources SMA-17, SMA-18, and SMA-19.

(^h) simultaneous fit of the group of sources SMA-21 and SMA-22.

ration values). Panel (b) shows that the distributions of the 1st nearest neighbor distribution (open, black line histogram) peaks at the median value of 6.5'' with a σ of 4.7''; the 7th nearest neighbor distribution (filled, black line histogram) peaks at the median value of 31.5'' with a σ of 8.8''².

In order to determine if these two distributions correspond to bona fide characteristic spatial scalings of the region, we carried out 10,000 Monte Carlo simulations of a randomly spatially distributed sample of sources, of the same number as our SMA detected sources, and located within an area of the same size. Panel (a) of Figure 4 shows how the median separations for this randomly spatial distributed sample (red) increases monotonically, presenting only one minima, for $N=1$ (we found that the $\sigma_{Nth\text{ NNS}}$ function for a random distribution always increases monotonically, irrespective of the number of sources or the size of the area in which they are distributed). The 1st NNS for the simulated sample has a median value of 14.0'' and σ of 8.6'', which is greater than the homologous measured values by a factor larger than 2. Panel (b) shows how the randomly spatial distributed sample presents broad and shallow peaks for their 1st

and 7th nearest neighbor distributions (red) that cannot describe the narrow, well defined, peaks of the observed distributions of the SMA sources. The two spatial scalings we measure are therefore not random and are a characteristic feature of OMC 1n. We also carried out a two-point correlation function and found positive correlation at the same two spatial scales (see Figure A.1). We continue our analysis assuming these spatial scalings are a result of the fragmentation of the parental structures that gave rise to the small-scale clumps and cores in OMC 1n.

4. Discussion

The size and density of the new SMA sources, combined with a general lack of counterpart infrared sources and the presence of CO outflows, indicate these compact SMA sources are either cores with young protostars in the Class 0 evolutionary phase or pre-stellar cores. From Takahashi et al. (2013) we know the compact submillimeter sources in OMC 3 range in masses from 0.3 M_{\odot} to 5.7 M_{\odot} , and we find in this work that the masses of the OMC 1n sources range between 0.1 M_{\odot} and 3.0 M_{\odot} . The OMC 3 projected source sizes range from 1400 au to 8200 au, whereas the projected OMC 1n source sizes range from 400 au to 1300 au. In terms of density, sources in OMC 3 have $2.0 \times 10^8 \text{ cm}^{-3} > n_{H_2} > 1.9 \times 10^6 \text{ cm}^{-3}$, whereas the OMC 1n sources are on average less denser, and have $2.6 \times 10^7 \text{ cm}^{-3} > n_{H_2} > 2.8 \times 10^6 \text{ cm}^{-3}$. Finally, comparison of

² For a sample of sources that excludes SMA-9, -12, and -18, we find that the 1st NNS has a median value of 7.6'' with a σ of 4.3'', and the 7th NNS has a median value of 38.6'' and a corresponding σ of 9.4''. Although the measured parameters of these sources have larger error bars (c.f. Table 3), they are fully consistent with the statistical trend of the other sources.

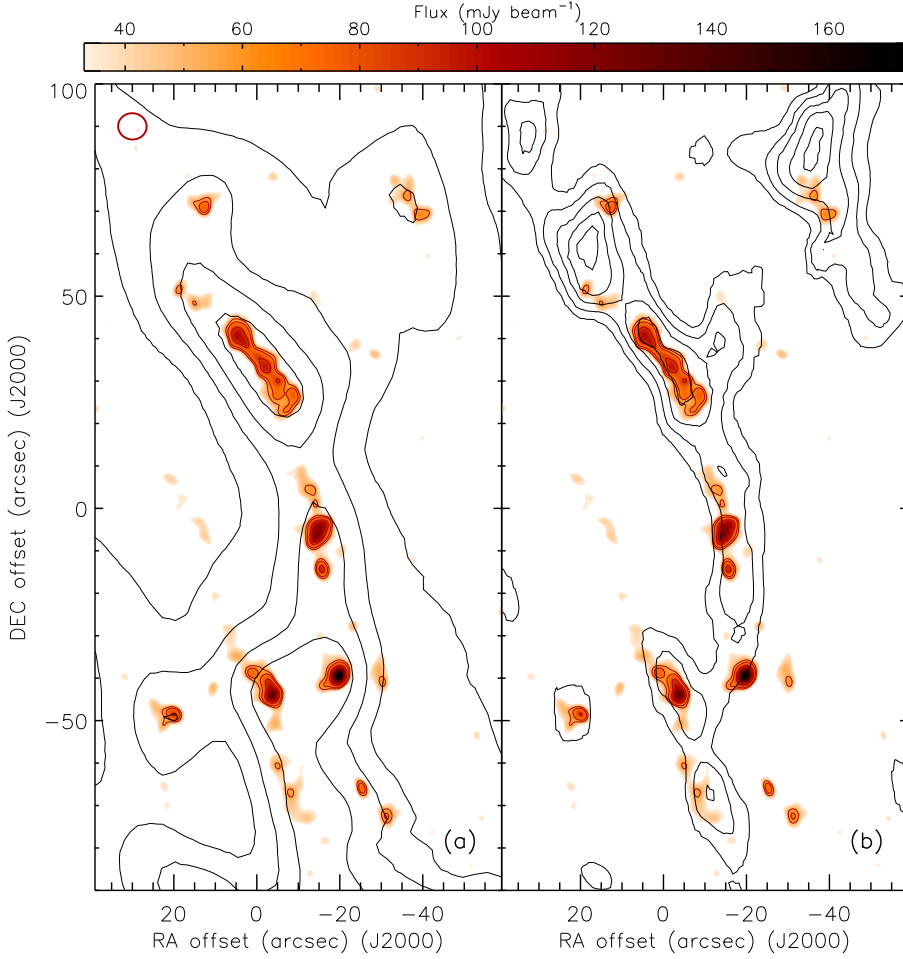


Fig. 3. Comparison of SMA continuum emission (grayscale) with: (a) 850 μm JCMT emission, where the JCMT contours range from 0.5 to 20 Jy beam^{-1} in steps of 1 Jy beam^{-1} (Johnstone & Bally 1999), and (b) VLA NH_3 (1,1) emission integrated from 4.6 to 12.3 km s^{-1} , where two lowest NH_3 contours are 20 and 100 $\text{mJy beam}^{-1} \text{km s}^{-1}$ and the higher level contours continue in steps of 100 $\text{mJy beam}^{-1} \text{km s}^{-1}$ (Wiseman & Ho 1998). Coordinate offsets are measured with respect to $(\alpha, \delta) = (05^{\text{h}}35^{\text{m}}17.0^{\text{s}}, -05^{\circ}22'00.0'')$ (J2000.0). [This figure is available in color in the online electronic journal.]

the fraction of infrared counterparts, 67% for OMC 3 and 8% for OMC 1n, indicates that the latter are less evolved than the former. We find that the OMC 1n sources are indeed very young objects that have not had enough time to substantially move away from their birth site within the filament. Class 0 protostars ($\lesssim 10^4$ yr; Larson 2003) with an average velocity of 1 km s^{-1} may travel up to ~ 0.01 pc away from their birth site during this evolutionary phase. As such, it is valid to infer the fragmentation scale of the parental filament from the measured spatial distribution of these very young sources.

The measured separations are projected distances. The observed scales, λ_{obs} , would therefore correspond to the fragmentation scales, λ_{frag} , if the inclination, i , of the filament and small-scale clumps with respect to the line-of-sight is 90° [$\lambda_{\text{obs}} = \lambda_{\text{frag}} \cdot \sin(i)$]. As discussed in section § 3.3 and shown in Figure 4, there are two observed spatial scales present in OMC 1n: a larger scale corresponding to $\lambda_{\text{obs}_1} = 31'' \pm 9''$ (0.06 ± 0.02 pc), and a smaller scale corresponding to $\lambda_{\text{obs}_2} = 6'' \pm 5''$ (0.01 ± 0.01 pc). The λ_{obs_1} scale corresponds to the distance between the JCMT/SCUBA dust emission small-scale clumps and to the distance between the VLA NH_3 small-scale clumps (see Figure 3). The λ_{obs_2} scale corresponds to the separation of the new SMA compact sources within these small-scale clumps. The inclination of the OMC 1n filament is unknown. The median inclination of a randomly oriented filament is 60° (Hanawa et al. 1993); using this angle of inclination the fragmentation scales would be $\lambda_{\text{frag}_1} = 36''$ (0.07 pc) and $\lambda_{\text{frag}_2} = 7''$ (0.01 pc). The projection effect is thus smaller than the error in λ_{obs_1} and λ_{obs_2}

for an inclination of 60° . In fact, the projection effect is only significant if the inclination angle is smaller than 50° .

A filament may be supported against gravitational collapse and fragmentation through different physical means, of which we consider the following three: thermal pressure, turbulent pressure, and magnetic pressure. These physical processes may affect differently the final fragmentation scale of the filament. In addition, the kinematics of the filament and small-scale clumps may also imprint on the spatial distribution of the sources formed therein. The filament could be experiencing large-scale collapse. Within the filament, the small-scale clumps could also be undergoing local collapse. The large scale collapse of the filament may shorten the distance between the small-scale clumps, and likewise, local collapse of the small-scale clumps may shorten the distance between the cores therein. The measured fragmentation scales may thus give us insight into which of these aforementioned processes are dominant. We discuss briefly each of these scenarios.

Let us first consider at what spatial scales the filament would fragment if no support against gravitational collapse is provided by either turbulence or magnetic fields, i.e., if we were to consider only thermal pressure. The Jeans length (Spitzer 1998), λ_{Jeans} , given by:

$$\lambda_{\text{Jeans}} = \sqrt{\frac{\pi c_s^2}{G \rho_0}} \quad (2)$$

where c_s is the sound speed given by $\sqrt{(k_B T)/(\mu m_H)}$ (k_B is the Boltzmann constant, T is the temperature, μ is the mean

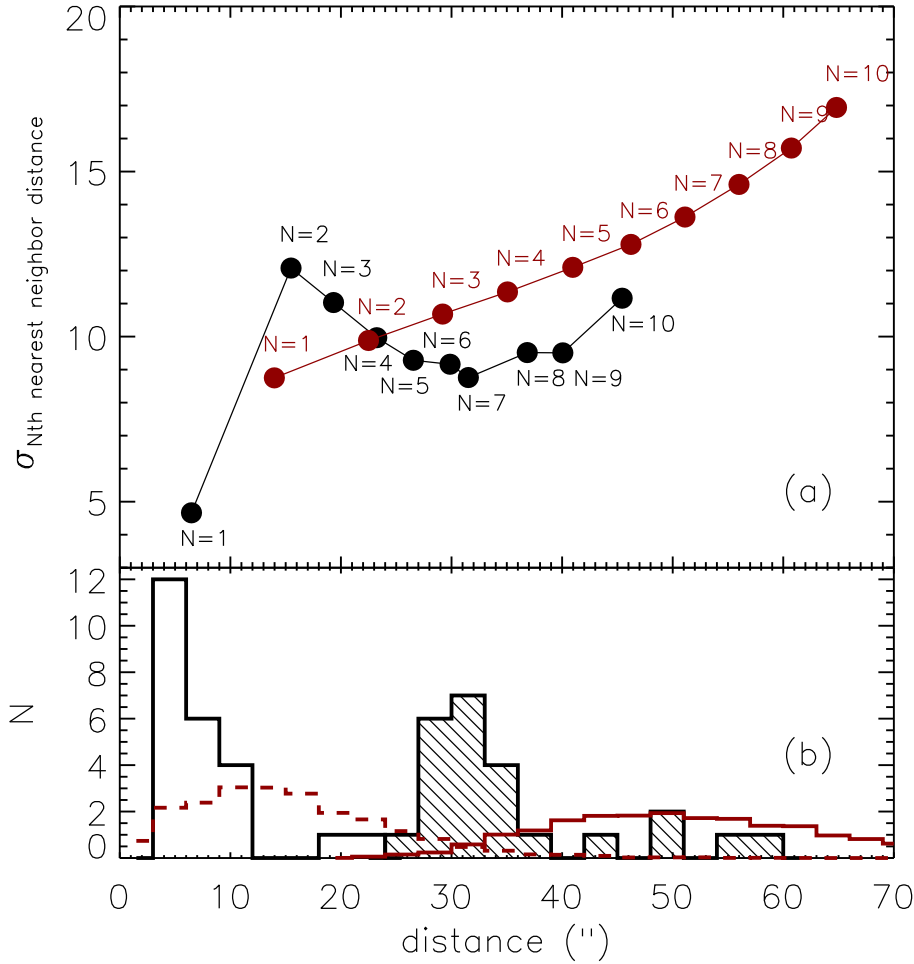


Fig. 4. (a) Median Nth nearest neighbor separations of the SMA sources vs. σ_{Nth} NNS, where each data point (black filled symbol) is labeled from $N = 1, \dots, 10$. The red symbols represent the median Nth NNS for a random spatial distribution for comparison purposes (see text). (b) Distribution of the 1st NNS (open black histogram), and the distribution of the 7th NNS (filled black histogram). The red dashed-line and the red solid-line histograms correspond to the 1st and 7th NNS for random spatial distributions, respectively. All the histogram bin widths correspond to the approximate SMA synthesized beams size, $3''$.

molecular weight, and m_H is the proton mass), G is the gravitational constant, and ρ_0 is the initial volume density given by $n_0/(\mu m_H)$. We calculated the Jeans length of the filament, $\lambda_{Jeans,1}$, by using the peak density of the OMC 1n filament ($n_{H_2} \sim 10^5 \text{ cm}^{-3}$, Johnstone & Bally 1999) and a temperature of 15 K (Wiseman & Ho 1998). For these conditions the local Jeans length for the filament is 0.08 pc ($40''$). For a temperature of 10 K, $\lambda_{Jeans,1} = 0.07 \text{ pc}$ ($35''$). The larger fragmentation scale measured $\lambda_{obs,1}$ is thus consistent with $\lambda_{Jeans,1}$ for inclinations angles greater than 50° . It is therefore possible that the OMC 1n filament underwent thermal fragmentation giving origin to the small-scale clumps therein. For the analysis of the fragmentation of the small-scale clumps themselves, we calculated the local Jeans length using Equation 2, a density of $5 \times 10^5 \text{ cm}^{-3}$ (Wiseman & Ho 1998), and a temperature of 15 K as before: $\lambda_{Jeans,2} = 0.033 \text{ pc}$ ($16''$). For a temperature of 10 K (canonical value of the temperature of a dense pre-stellar region, i.e., that has not yet undergone internal heating), $\lambda_{Jeans,2} = 0.026 \text{ pc}$ ($13''$). The measured separation between individual cores within the small-scale clumps $\lambda_{obs,2}$ is smaller than $\lambda_{Jeans,2}$ by a factor of ~ 2 -3. The measured spatial lengths and Jeans lengths are summarized in Table 5.

For comparison, our previous SMA study of the OMC 3 region in the \int -shaped filament (Takahashi et al. 2013) detected 12 spatially resolved continuum sources at $850 \mu\text{m}$. We measured a quasi-periodic separation between these individual sources of $\sim 17''$ (0.035 pc). This spatial scale is smaller than the local thermal Jeans length by a factor of 2.0-7.3.

Our finding – *2 level hierarchical fragmentation in OMC 1n* – is in general agreement with studies in other star forming regions, notably in the Spokes cluster in NGC 2264 (Teixeira et al. 2006, 2007), in the more massive infrared dark cloud IRDC G11.11-0.12 (Kainulainen et al. 2013), and in other massive star forming regions (Palau et al. 2015). Further observational evidence for this type of fragmentation is given in a recent study of the young stellar population of the Orion A and B clouds, where Megeath et al. (2015) find that the median separation between sources is similar to the Jeans length. In all these aforementioned regions, for spatial scales smaller than 0.5 pc, the respective fragmentation lengths observed are consistent with the local spherical Jeans length. For spatial scales larger than 0.5 pc, Kainulainen et al. (2013) find that the observed fragmentation scale is consistent with the prediction from the gravitational instability of a self-gravitating isothermal cylinder (e.g. Nagasawa 1987). We explore next what physical processes could explain the smaller fragmentation length measured in OMC 1n from the separations of the individual cores.

Turbulence will produce a fragmentation scale larger than the Jeans length, as discussed in Takahashi et al. (2013). Theoretical fragmentation length scales are directly proportional to the sound speed, c_s . Turbulence can be included in this calculation by using an effective sound speed, c_{eff} , instead of c_s ; c_{eff} is given by $\sqrt{(\Delta v_t^2 + \Delta v_{nt}^2)/8 \ln 2}$, where Δv_t and Δv_{nt} correspond to the thermal and non-thermal velocity linewidths, respectively. Turbulence is included in the Δv_{nt} term. In the presence of turbulence, $c_{eff} > c_s$ and the fragmentation scale is conse-

Table 4. Calculated properties of the SMA sources and their nearest neighbor separation.

ID	R _{geom} (AU)	Mass ^a (M _⊙)	\bar{n}_{H_2} (10 ⁷ cm ⁻³)	NNS (arcsec)
SMA-1	661 ± 195	0.5 ± 0.2	6.6 ± 3.4	5.8 ± 0.6
SMA-2	1012 ± 234	1.0 ± 0.3	3.5 ± 1.4	5.8 ± 0.6
SMA-3	949 ± 258	1.0 ± 0.4	4.3 ± 2.0	20.5 ± 0.5
SMA-4	557 ± 177	0.5 ± 0.2	10.5 ± 5.9	4.9 ± 0.6
SMA-5	436 ± 243	0.4 ± 0.3	16.3 ± 15.8	4.9 ± 0.6
SMA-6	1285 ± 124	2.2 ± 0.3	3.7 ± 0.7	9.5 ± 0.4
SMA-7	1178 ± 210	1.7 ± 0.6	3.9 ± 1.5	6.9 ± 0.5
SMA-8	1217 ± 395	1.3 ± 0.8	2.6 ± 1.8	4.2 ± 0.7
SMA-9	594 ± 343	0.1 ± 0.1	2.3 ± 2.8	4.2 ± 0.7
SMA-10	906 ± 368	0.9 ± 0.6	4.3 ± 3.3	4.2 ± 0.7
SMA-11	1024 ± 292	0.9 ± 0.4	3.1 ± 1.6	4.0 ± 0.8
SMA-12	564 ± 407	0.1 ± 0.1	2.0 ± 3.0	4.0 ± 0.8
SMA-13	1151 ± 63	2.3 ± 0.2	5.4 ± 0.6	6.5 ± 0.6
SMA-14	629 ± 79	0.8 ± 0.1	11.1 ± 2.6	9.0 ± 0.3
SMA-15	570 ± 18	1.3 ± 0.1	26.4 ± 1.5	10.6 ± 0.4
SMA-16	757 ± 199	0.7 ± 0.2	5.5 ± 2.5	10.6 ± 0.4
SMA-17	534 ± 109	1.0 ± 0.3	22.6 ± 8.4	3.0 ± 0.5
SMA-18	694 ± 587	0.8 ± 1.6	9.0 ± 18.9	3.0 ± 0.5
SMA-19	917 ± 470	0.9 ± 0.8	4.2 ± 4.2	4.7 ± 0.4
SMA-20	771 ± 124	0.9 ± 0.2	7.3 ± 2.1	23.6 ± 0.5
SMA-21	939 ± 417	0.8 ± 0.5	3.3 ± 2.6	7.2 ± 0.5
SMA-22	810 ± 249	0.6 ± 0.3	4.3 ± 2.3	7.2 ± 0.5
SMA-23	850 ± 89	0.5 ± 0.1	3.2 ± 0.6	8.8 ± 0.5
SMA-24	969 ± 116	0.9 ± 0.2	3.6 ± 0.8	8.8 ± 0.5

Notes. ^(a) If a warmer T_d of 20 K is used, then the masses of the sources are lower by a factor of ~ 1.5 . Taking into account the estimated thermal free-free emission from jets, the masses listed may be lower by a value smaller than 0.1 M_⊙.

Table 5. Summary of measured spatial scales and corresponding local Jeans lengths.

Parental Structure	Measured Scale	Local Jeans Length
filament	$\lambda_{obs,1} = 31'' \pm 9''$	$\lambda_{Jeans,1}(10\text{ K}) = 35''(0.07\text{ pc})$ $\lambda_{Jeans,1}(15\text{ K}) = 40''(0.08\text{ pc})$
small-scale clump	$\lambda_{obs,2} = 6'' \pm 5''$	$\lambda_{Jeans,2}(10\text{ K}) = 13''(0.026\text{ pc})$ $\lambda_{Jeans,2}(15\text{ K}) = 16''(0.033\text{ pc})$

quently larger than the thermal Jeans length. Since the two measured fragmentation scales, $\lambda_{obs,1}$ and $\lambda_{obs,2}$, are either equal or smaller than their respective local Jeans lengths we can rule out turbulence pressure as a dominant physical process in the fragmentation of the filament and small-scale clumps, and conclude that it has essentially decayed for these size scales (< 0.3 pc) in OMC 1n. This conclusion is consistent with that previously reached by Houde et al. (2009), where they find turbulence does not dominate the dynamics in OMC 1. Takahashi et al. (2013) also rule out turbulence as a dominant physical process in the fragmentation of OMC 3 and further suggest that it has dissipated for size scales ≤ 1 pc.

We discuss next the role the magnetic field may play in the fragmentation of OMC 1n. Several polarimetric observations indicate that the magnetic field in OMC 1n is mostly perpendicular to the filament and its small-scale clumps, namely, far-infrared observations at 100 μ m from the Kuiper Airborne Observatory (Schleuning 1998), submillimeter observations from CSO at 350 μ m (Schleuning 1998; Houde et al. 2004) and JCMT observations at 760 μ m (Vallée & Bastien 1999). The orientation of

the magnetic field relative to the axis of the filamentary cloud has important implications in terms of its stability (e.g. Nagasawa 1987). Stodólkiewicz (1963) found that the critical length of perturbation increases with increasing magnetic field intensity, for the idealized case of a magnetic field perpendicular to an infinitely long cylinder of compressible gas. Assuming the orientation of the larger magnetic field is preserved for smaller scales in the filament (Li et al. 2009), and that the intensity of the magnetic field increases for smaller scales, it follows that the magnetic field topology of OMC 1n cannot explain why $\lambda_{obs,2}$ is smaller than $\lambda_{Jeans,2}$ (Table 5). Regarding the strength of the magnetic field, B , Houde et al. (2004) calculate a total magnitude of $B \approx 850 \mu$ G, using the observed line-of-sight magnetic field strength of 360 μ G determined from Zeeman detection in CN Crutcher (1999) and an inclination of 65° (consistent with a plane-of-the sky magnetic field strength of 760 μ G determined from polarimetric observations (Houde et al. 2009)). They conclude that OMC 1 is magnetically supercritical. The mass-to-flux ratio in units of the critical mass-to-flux ratio, γ , (Crutcher 2004) is given by:

$$\gamma = \frac{(M/\Phi)_{observed}}{(M/\Phi)_{critical}} = 7.6 \times 10^{-21} \frac{N(H_2)}{B}. \quad (3)$$

Using the aforementioned value of B , and an $N(H_2)$ of $5 \times 10^{23} \text{ cm}^{-2}$ obtained from Herschel observations (Lombardi et al. 2014), we calculate γ to be ≈ 4.5 , and confirm that OMC 1n is magnetically supercritical. Houde et al. (2004) also calculate a magnetic-to-gravitational energy ratio of ≈ 0.26 . Our current data do not allow us to analyse the magnetic field properties in more detail, however, from the above arguments it appears that the magnetic field does not provide enough support against gravitational collapse, nor does it explain the shorter separations between cores observed in the small scale clumps.

Having ruled out turbulence and the magnetic field as responsible agents for the smaller fragmentation scales measured in OMC 1n, $\lambda_{obs,2}$, we now explore other possible processes. As already discussed, thermal fragmentation alone cannot fully explain $\lambda_{obs,2}$, however, if thermal fragmentation were to occur concomitantly with the local collapse of the small-scale clumps, naturally the separations between the individual sources formed within the small-scale clumps are expected to decrease. Pon et al. (2011) show that local collapse proceeds on a much smaller timescale than global collapse for filamentary structures. This analytical model would be consistent with the scenario where the local collapse in the small-scale clumps is occurring at a much faster pace than the global collapse of OMC 1n, thus the shortening of the distances between sources would be more pronounced within the small-scale clumps, rather than in between them. We propose this may indeed be an important mechanism at play in OMC 1n. We plan to explore this scenario and constrain the effect that local vs. global collapse has on the final spatial separations of the sources with future ALMA data.

5. Summary

1. We found 24 new compact millimeter sources in OMC 1n that are likely young protostars in the Class 0 evolutionary phase, or pre-stellar objects.
2. The OMC 1n sources range in size from 400 au to 1300 au, and range in mass from 0.1 M_⊙ to 2.3 M_⊙ (assuming a dust temperature of 15 K and a distance to OMC 1n of 414 pc). These measured values are lower limits as our data is missing the zero spacing flux emitted by more extended structures.

3. The spatial distribution of the sources reveal a hierarchical two-level fragmentation the parental filamentary structure - the larger scale, $\lambda_{\text{obs},1} = 31'' \pm 9'' (0.06 \pm 0.02 \text{ pc})$, is associated with the fragmentation of the filament into small-scale clumps, while the smaller scale, $\lambda_{\text{obs},2} = 6'' \pm 5'' (0.01 \pm 0.01 \text{ pc})$, is associated with the fragmentation of the small-scale clumps into cores/SMA sources.
4. Comparison of these two measured scales with the local Jeans lengths of the parental filament and of the small-scale clumps shows that the larger observed scale, $\lambda_{\text{obs},1}$ is consistent with thermal fragmentation for inclinations greater than 50° . The smaller scale, $\lambda_{\text{obs},2}$, is smaller than the local Jeans lengths of the parental small-scale clumps by a factor of ~ 2 -3.
5. We discuss other mechanisms that may explain the smaller measured scale, $\lambda_{\text{obs},2}$. Further molecular line data is needed to disentangle the different physical processes involved, such as local collapse within OMC 1n.

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References

- André, P., Men'shchikov, A., Bontemps, S., et al. 2010, *Astronomy and Astrophysics*, 518, L102
- Anglada, G., Villuendas, E., Estalella, R., et al. 1998, *AJ*, 116, 2953
- Bally, J., Lanber, W. D., Stark, A. A., & Wilson, R. W. 1987, *Astrophysical Journal*, 312, L45
- Bally, J., Testi, L., Sargent, A., & Carlstrom, J. 1998, *AJ*, 116, 854
- Beckwith, S. V. W., Sargent, A. I., Chini, R. S., & Guesten, R. 1990, *AJ*, 99, 924
- Crutcher, R. M. 1999, *ApJ*, 520, 706
- Crutcher, R. M. 2004, *The Magnetized Interstellar Medium*, 123
- Hacar, A., Tafalla, M., Kauffmann, J., & Kovács, A. 2013, *Astronomy and Astrophysics*, 554, A55
- Hanawa, T., Nakamura, F., Matsumoto, T., et al. 1993, *Astrophysical Journal*, 404, L83
- Heitsch, F. 2013, *Astrophysical Journal*, 769, 115
- Ho, P. T. P., Moran, J. M., & Lo, K. Y. 2004, *Astrophysical Journal*, 616, L1
- Houde, M., Dowell, C. D., Hildebrand, R. H., et al. 2004, *Astrophysical Journal*, 604, 717
- Houde, M., Vaillancourt, J. E., Hildebrand, R. H., Chitsazzadeh, S., & Kirby, L. 2009, *Astrophysical Journal*, 706, 1504
- Ivezić, Ž., Connolly, A., VanderPlas, J., & Gray, A. 2013, *Statistics, Data Mining, and Machine Learning in Astronomy*, by Ž. Ivezić et al. Princeton University Press, 2013,
- Johnstone, D. & Bally, J. 1999, *Astrophysical Journal*, 510, L49
- Kainulainen, J., Ragan, S. E., Henning, T., & Stutz, A. 2013, *A&A*, 557, A120
- Kirk, H., Myers, P. C., Bourke, T. L., et al. 2013, *Astrophysical Journal*, 766, 115
- Landy, S. D., & Szalay, A. S. 1993, *ApJ*, 412, 64
- Larson, R. B. 2003, *Reports on Progress in Physics*, 66, 1651
- Li, H.-b., Dowell, C. D., Goodman, A., Hildebrand, R., & Novak, G. 2009, *ApJ*, 704, 891
- Lombardi, M., Bouy, H., Alves, J., & Lada, C. J. 2014, *Astronomy and Astrophysics*, 566, A45
- Megeath, S. T., Gutermuth, R., Muzerolle, J., et al. 2015, *arXiv:1511.01202*
- Menten, K. M., Reid, M. J., Forbrich, J., & Brunthaler, A. 2007, *Astronomy and Astrophysics (ISSN 0004-6361)*, 474, 515
- Muench, A., Getman, K., Hillenbrand, L., & Preibisch, T. 2008, *Handbook of Star Forming Regions*, 483
- Mundy, L. G., Looney, L. W., & Lada, E. A. 1995, *ApJ*, 452, L137

- Myers, P. C. 2009, *Astrophysical Journal*, 700, 1609
- Myers, P. C. 2011, *The Astrophysical Journal*, 735, 82
- Nagasawa, M. 1987, *Progress of Theoretical Physics*, 77, 635
- O'Dell, C. R. 2001, *Annual Review of Astronomy and Astrophysics*, 39, 99
- O'Dell, C. R., Muench, A., Smith, N., & Zapata, L. 2008, *Handbook of Star Forming Regions*, 544
- Ostriker, J. 1964, *Astrophysical Journal*, 140, 1056
- Palau, A., Ballesteros-Paredes, J., Vázquez-Semadeni, E., et al. 2015, *MNRAS*, 453, 3785
- Panagia, N., & Felli, M. 1975, *A&A*, 39, 1
- Peterson, D. E. & Megeath, S. T. 2008, *Handbook of Star Forming Regions*, 590
- Pineda, J. E. & Teixeira, P. S. 2013, *Astronomy and Astrophysics*, 555, A106
- Pon, A., Johnstone, D., & Heitsch, F. 2011, *Astrophysical Journal*, 740, 88
- Reipurth, B., Rodríguez, L. F., & Chini, R. 1999, *AJ*, 118, 983
- Sault, R. J., Teuben, P. J., & Wright, M. C. H. 1995, *Astronomical Data Analysis Software and Systems IV*, 77, 433
- Schleuning, D. A. 1998, *Astrophysical Journal*, 493, 811
- Schneider, N., Csengeri, T., Bontemps, S., et al. 2010, *Astronomy and Astrophysics*, 520, A49
- Schneider, S. & Elmegreen, B. G. 1979, *Astrophysical Journal Supplement Series*, 41, 87
- Scoville, N. Z., Carlstrom, J. E., Chandler, C. J., et al. 1993, *PASP*, 105, 1482
- Spitzer, L. 1998, *Physical Processes in the Interstellar Medium*, by Lyman Spitzer, pp. 335. ISBN 0-471-29335-0. Wiley-VCH, May 1998.,
- Stodólkiewicz, J. S. 1963, *Acta Astronomica*, 13, 30
- Tackenberg, J., Beuther, H., Henning, T., et al. 2014, *Astronomy and Astrophysics*, 565, A101
- Takahashi, S., Ho, P. T. P., Teixeira, P. S., Zapata, L. A., & Su, Y.-N. 2013, *Astrophysical Journal*, 763, 57
- Takahashi, S., Teixeira, P. S., Ho, P. T. P., & Zapata, L. A. 2015, in prep.
- Tatematsu, K., Kandori, R., Umemoto, T., & Sekimoto, Y. 2008, *Publications of the Astronomical Society of Japan*, 60, 407
- Teixeira, P. S., Lada, C. J., Young, E. T., et al. 2006, *Astrophysical Journal*, 636, L45
- Teixeira, P. S., Zapata, L. A., & Lada, C. J. 2007, *ApJ*, 667, L179
- Vallée, J. P. & Bastien, P. 1999, *Astrophysical Journal*, 526, 819
- VanderPlas, J., Connolly, A. J., Ivezić, Z., & Gray, A. 2012, *Proceedings of Conference on Intelligent Data Understanding (CIDU)*, pp. 47-54, 2012., 47
- Wilner, D. J. & Welch, W. J. 1994, *Astrophysical Journal*, 427, 898
- Wiseman, J. J. & Ho, P. T. P. 1998, *Astrophysical Journal* v.502, 502, 676
- Zapata, L. A., Rodríguez, L. F., Kurtz, S. E., O'Dell, C. R., & Ho, P. T. P. 2004, *ApJ*, 610, L121

Appendix A: Additional spatial correlation analysis

We also performed a 2-point correlation analysis of the spatial distribution of the SMA sources, using the Landy & Szalay (1993) estimator. The correlation function, $\omega(\theta)$, is therefore given by:

$$\omega(\theta) = \frac{DD(\theta) - 2 \cdot DR(\theta) + RR(\theta)}{RR(\theta)}, \quad (\text{A.1})$$

where $DD(\theta)$ corresponds to the distribution of measured angular separations, θ , between pairs of SMA sources. $RR(\theta)$ corresponds to the distribution of angular separations between pairs of random sources, and $DR(\theta)$ corresponds to the distribution of angular separations between pairs of an SMA source and a random source. The sample of random sources consisted of the same number of sources as the sample of SMA sources, placed within the same RA and DEC ranges, and was generated 10 000 times; the final $RR(\theta)$ and $DR(\theta)$ distributions were averaged along θ .

Figure A.1 shows that $\omega(\theta)$ has two maxima, denoting preferred spatial correlations at separations consistent with $\lambda_{\text{obs},1}$ and $\lambda_{\text{obs},2}$, as measured above using the median 1st and 7th nearest neighbor separations of the SMA sources, respectively.

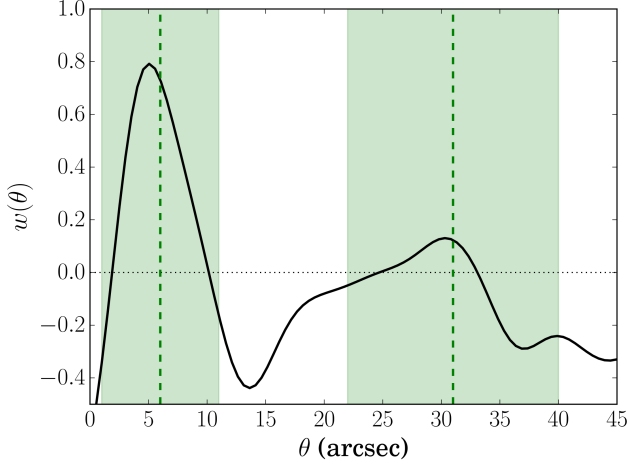


Fig. A.1. Correlation function, $\omega(\theta)$, of the angular separation, θ , between the SMA sources. The vertical dashed lines correspond to λ_{obs_1} and λ_{obs_2} , respectively. The width of the shaded areas correspond to the corresponding error of λ_{obs_1} and λ_{obs_2} . The dotted horizontal marks $\omega(\theta)=0$.